

On Mars' Atmospheric Sputtering after MAVEN first Martian year of

Measurements

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Abstract:

Mars likely lost a significant part of its atmosphere to space during its history. The sputtering of the atmosphere, by precipitating planetary heavy pickup ions accelerated by the solar wind, is one of the processes that could have significantly contributed to this atmospheric escape, in particular since the cessation of its global magnetic field, 4.0-4.1 Gyr ago. We present a two-year baseline analysis of MAVEN observations of the precipitating flux. We use this measurement to model the expected escape rate and exospheric structure induced by this precipitation. We conclude that sputtering signatures in the dayside exosphere will be difficult to identify by MAVEN and the induced atmospheric escape of O atoms remains orders of magnitude smaller than the expected rate induced by dissociative recombination of O_2^+ in Mars' ionosphere. On the contrary, deep in the nightside, Mars' sputtering might be the main source of the non-thermal part of the exospheric density profiles of species with mass larger or equal to Ar.

Plain Language Summary

After two years of MAVEN measurements, it is possible to reconstruct the average precipitating heavy ion flux into Mars' atmosphere. These ions can lead to the sputtering of Mars' atmosphere, a process which might have eroded a significant part of Mars' atmosphere 4 Gyr ago. Today, it is much more difficult to identify a clear signature of this process. However, we here show that such signature might be found deep in the night according to models.

I Introduction

The main objective of Mars Atmosphere and Volatile EvolutioN (MAVEN) mission is to characterize Mars' present atmospheric escape, its main drivers and to reconstruct its past evolution (Jakosky et al. 2015). To achieve these goals, MAVEN's strategy consists of identifying signatures of the various channels of atmospheric escape and estimating their importance and dependence with respect to solar wind conditions (Lillis et al. 2015). Among these channels, sputtering of the Martian atmosphere by precipitating heavy ions has been identified as one of these potential channels (Luhmann and Kozyra 1991). The first evidence of heavy ion precipitation for nominal solar wind conditions has been reported by Leblanc et al. (2015), highlighting that sputtering should clearly occur at Mars. Hara et al. (2017a), analyzing the first one year and a half of MAVEN measurements, showed that this precipitation is organized with respect to the orientation of the solar wind convection electric field and is dependent on the upstream solar wind conditions. Hara et al. (2017b) also highlighted the role of the crustal magnetic field in the local increase of this precipitating flux.

However, no clear signature of the consequences of these precipitating heavy ions on the structure of the Martian atmosphere could be identified, as an example, in the oxygen atomic exosphere or in the thermosphere (Fang et al. 2013; Leblanc et al. 2017). Bhardwaj et al. (2017) suggested that the part of the Ar density profiles measured by NGIMS and MENCA/MOM above 250 km in altitude could be significantly changed by such precipitation, a conclusion we will address here by modelling the expected density profiles associated with the typical range of ion precipitation measured by MAVEN.

In this paper, we present the first estimate of sputtered escape induced by the precipitating energy flux measured by MAVEN. We also determine how the exosphere could be shaped by the sputtered atmospheric ejecta by using the Exospheric Global Model (EGM; Leblanc et al. 2017) developed to describe Mars' exosphere at any Martian season and solar activity. We compare the spatial structure of the exosphere as produced by sputtering with the one produced by dissociative recombination of the dominant ion O_2^+ in Mars' ionosphere and estimate the contribution of sputtering to Mars'

atmospheric erosion. In section II, we present MAVEN measurements of precipitating ions based on one Martian year of measurements. In section III, the simulated signatures induced by these precipitating ions are described using EGM and in section IV we conclude.

II Measurements

In order to reconstruct the flux of precipitating ion measured by MAVEN, we used the same approach as in Leblanc et al. (2015) and Leblanc et al. (2017), that is, we used all available measurements of the ion mass and energy distributions realized between 200 and 350 km in altitude by the Solar Wind Ion Analyzer (SWIA; Halekas et al. 2015) and by the Supra-Thermal And Thermal Ion Composition (STATIC; McFadden et al. 2015). STATIC is an energy, mass and angular ion spectrometer, operating in the range of 0.1 eV/q to 30 keV/q (energy resolution between 11 and 16%), with a Field-of-View (FOV) of 360°x90° (with 22.5° resolution) and a mass range from 1 to 70 amu. STATIC has a mass resolution large enough to separate the main ion species: O₂⁺, CO₂⁺, and O⁺. For this study, we used the d0 product of STATIC measurements with 128 to 512 second time resolution corresponding to 32 bins in energy, 8 bins in mass and 16 bins along the polar direction. STATIC is mounted on the Articulated Payload Platform (APP), which points STATIC's field of view into the RAM direction at periapsis. SWIA is an energy and angular ion spectrometer covering an energy range 25 eV/q to 25 keV/q (with 48 logarithmically spaced energy steps) and a broad 360°x90° FOV. SWIA is positioned on the spacecraft (S/C) deck in order to observe the direction of the solar wind most of the time but also to observe the nadir direction around periapsis. In this paper, we used the coarse survey (cs) product of SWIA with an angular coverage of 360°x90°, an angular resolution of 22.5°x22.5° and a temporal resolution of 4 seconds.

In order to reconstruct the energy distribution of the precipitating flux, we selected all portions of MAVEN orbit between 200 and 350 km in altitude. We then used data from anodes whose field of view (FOV) was within the 75° cone angle centered along the zenith direction (Leblanc et al. 2015).

We then selected the measurements during which the average FOV of each instrument during the path between 200 and 350 km covered more than 70% of this 75° cone angle. This criterion restricts us to a relatively small fraction of the observations performed by MAVEN during this Martian year (between the 12/2/2014 and the 10/31/2016), namely:

- for SWIA, 785 paths on the nightside and 839 paths on the dayside, from the 7244 available measurements,
- for STATIC d0, 124 paths on the nightside and 259 paths on the dayside, from the 7268 available measurements.

The global coverage of these set of measurements still allows us to obtain uniform coverage of the MSO latitude-longitude frame. Figure 1a displays the average precipitating flux measured by MAVEN SWIA and STATIC during the first two years of MAVEN operations (between 12/02/2014, Ls=243°, and 10/31/2016, Ls=247°).

FIGURE 1

As shown in Figure 1a, there is globally a good agreement (within the dispersion of the measurements indicated by the vertical bars) between the averaged measured flux by SWIA (red line) and STATIC (blue and green lines), despite the fact that SWIA field of view allows a better coverage of the cone angle pointing towards the zenith than STATIC. The comparison suggests that, when averaging the day and night sides, there is an equivalent number of O⁺ (green line) and H⁺ (blue line) precipitating towards the atmosphere. In Figure 1b, we plotted the reconstructed precipitating fluxes for the measurements done by MAVEN on the dayside (solid line) and on the nightside (dashed line). There is an asymmetry between day and night sides, with a flux approximately two times larger on the dayside with respect to the nightside above 100 eV (red solid and dashed lines). Comparing STATIC flux distribution for masses between 12 and 21 amu (green dashed lines) and SWIA flux (all masses, dashed red line), the precipitating flux on the nightside

appears to be dominated below 1000 eV by ions with masses other than the O⁺ mass. Such results were also suggested by Hara et al. (2017).

III Application to EGM

Using the precipitating flux measured by MAVEN for the day and night sides (STATIC d0 product measured fluxes displayed with green lines in Figure 1b), we simulated the fate of precipitating pick-up O⁺ ions into Mars' atmosphere using Exospheric Global Model (EGM) (Leblanc et al. 2017) coupled to Laboratoire de Météorologie Dynamique-Global Circulation Model (LMD-GCM), including the ion dynamics (Chaufray et al. 2014) in order to calculate the induced exosphere and atmospheric escape by the precipitating ion. EGM is a multispecies parallelized 3D model describing, for any season and solar conditions, how energetic O⁺ ions collisionally interact with Mars' upper atmosphere (simulated by LMD-GCM for the main neutral atmospheric species CO₂, CO, Ar, O and N₂). To describe the collision, we used at low energy the universal collision scheme of Lewkow and Karchenkov (2014) and at high energy molecular dynamics scheme (Leblanc and Johnson 2002) in order to take into account the possibility of molecular dissociation (Leblanc et al. 2017). Here, the LMD-GCM model provided a description of the state of the upper atmosphere, that is, the thermosphere and ionosphere, at Ls = 0°, 90°, 180° and 270° for solar mean conditions, representative of the period between 12/02/2014 and 10/24/2016.

In the following, we use EGM to describe the effect of these precipitating particles on Mars' atmosphere focusing on the induced exosphere and atmospheric escape. In order to compare these effects to other atmospheric sources of the exosphere or atmospheric escape, EGM is also used to describe the contribution of dissociative recombination of O₂⁺ to the exosphere and atmospheric escape, the 3D spatial distribution of the O₂⁺ being provided by LMD-GCM (Chaufray et al. 2014). We also include a description of the exospheric thermal component, that is, the thermal extension of any neutral atmosphere above the exobase (Yagi et al. 2012). In Figure 2, we shows profiles for the

three main species of Mars' atmosphere, O, CO₂ and N₂ as well as for Ar which is particularly interesting in the context of MAVEN observations and of the sputtering. Indeed, Jakosky et al. (2017) suggested that the Ar isotopic fractionation measured by MAVEN might be interpreted as a signature of Mars' past atmospheric escape due to sputtering.

FIGURE 2

As shown in Figure 2, there is a clear transition between an exosphere dominated by the thermal component (characterized by small scale height) and by the non-thermal component (characterized by large scale height and produced by dissociative recombination and sputtering). This transition depends on species and local time. As an example, it occurs around 600 km for the atomic oxygen exospheric component on the dayside (panels a, b, c, e and f) and around 500 km in altitude on the nightside (panel d), the thermosphere being colder on the nightside. For N₂, such a transition occurs slightly above 400 km, whereas it is around 350 km for CO₂ and Ar. There is a clear difference at low altitudes between CO₂ and Ar at the dawn equator due to the thermal component of the exosphere derived from the thermosphere modeled by LMD-GCM. The exospheric density profiles of CO₂ and Ar are close to each other on the nightside, at dawn and at the poles. They diverge gradually on the day side with increasing local time, the CO₂ exospheric density being up to 8 times larger at the dusk equator than the Ar density at 1500 km in altitude. This divergence is due to a change of the mixing ratio of Ar and CO₂ with increasing local time. These results displayed for Ls=180° in Figure 2 are essentially similar at the three other Ls considered in this work, Ls=0°, 90° and 270°.

FIGURE 3

Because the profiles in Figure 2 are built from three sources: a component from the thermal expansion of the atmosphere, a component induced by dissociative recombination of O_2^+ and one by sputtering, we separate these contributions in Figure 3 for two of the profiles displayed in Figure 2.

As illustrated in Figure 3 for O and Ar, the dissociative recombination of O_2^+ appears to be the main source of the exosphere for all species ejected into the exosphere on the dayside above 600 km in altitude (by two orders of magnitude with respect to the density induced by sputtering) and also for O, C, CO, N and N_2 on the nightside. In the case of Ar and CO_2 on the nightside, sputtering above 300 km in altitude is the main source of the exosphere. Dissociative recombination of O_2^+ is indeed not only a source of atomic oxygen escape and of a non-thermal oxygen component in the exosphere.

The dissociative recombination of O_2^+ produces non-thermal oxygen atoms with energies up to 3.5 eV which can collide with all thermospheric species. These collisions can therefore produce recoiled particles with energy significantly larger than the local thermal energy, recoiled particles that can reach high altitudes and populate the exosphere. Above a given altitude, these non-thermal particles get much more than the thermal exospheric particles leading to the two slopes density profile shown in Figure 3. Figure 3 shows that even for a species like Ar, the main origin of its density profile in the exosphere is not sputtering, as interpreted by Bhardwaj et al. (2017) but collisions with the non-thermal oxygen atoms produced by dissociative recombination. However, this conclusion is valid only on the dayside and on the nightside for some species. Indeed, species as O, C, N, CO and N_2 when ejected from the dayside into the exosphere can ballistically populate the whole nightside exosphere. But heavier species as CO_2 or Ar do not get enough energy on the dayside to reach the deep nightside. As a consequence, near midnight, the main source of exospheric CO_2 and Ar is sputtering, dissociative recombination being significantly reduced due to the very low O_2^+ density at such local time (few orders smaller than on the dayside; Chaufray et al. 2014).

FIGURE 4

In Figure 4, we displayed the evolution of the escape rate over one Martian year for both non-thermal processes considered in this work, the dissociative recombination of O_2^+ (panel a) and the sputtering (panel b). It might be surprising to find a non-zero escape rate of N_2 and CO induced by the dissociative recombination of O_2^+ , the escape energy of N_2 and CO being very close to the maximum energy released by the dissociative recombination of O_2^+ . Actually, we found that around 0.01% of the test-particles of N_2 or CO could reach 3 Mars radii (RM) which is the upper limit of our simulation domain where we considered particles as lost because having a large probability of ionization before reimpacting Mars. Indeed, only 2.3 eV are needed for a N_2 or CO molecule to reach 3 RM starting from the ionosphere, so that there is a small non-zero probability for a non-thermal O atom to provide such an energy to N_2 or CO. Averaging the values displayed in Figure 4 along one Martian year, the escape rates induced by dissociative recombination of O_2^+ is 1.5×10^{26} O/s, 2.7×10^{23} N_2 /s and 7.1×10^{22} CO/s. Supposing that the precipitating flux displayed in Figure 1 is representative of the average precipitation over a Martian year and that EGM provides a realistic description of sputtering, this process should then have induced between 12/2014 and 10/2016 an escape of 4×10^{23} O/s, 1.3×10^{23} C/s, 5.4×10^{23} N/s, 4.7×10^{23} CO/s, 1.2×10^{23} N_2 /s, 4.2×10^{22} Ar/s and 9.6×10^{22} CO_2 /s. Wang et al. (2015) calculated a theoretical analytical relation relating precipitating flux and induced atmospheric escape. Based on their approach, we would have found an escape rate of 1×10^{24} O/s (Leblanc et al. 2015) which is consistent with the smaller escape rate calculated in this paper, since a large part of the energy of the incident particle went into N and N_2 atmospheric loss rather than O, N and N_2 and was not considered in the Wang et al. (2015) simulation. We did not include other photo-chemical reactions that should also contribute to atmospheric escape and to the exosphere, namely, the dissociative recombination of N_2^+ (Fox and Hać 1997), the photo-dissociation of CO and dissociative recombination of CO^+ (Fox and Bakalian 2001) and the dissociative recombination of CO_2^+ (Fox 2004; Gröller et al. 2014). All these reactions are indeed

usually considered as marginally contributing to the atmospheric escape compared to O_2^+ dissociative recombination.

Lee et al. (2015) published estimates of the escape rate for solar mean conditions at the same four seasons as displayed in Figure 4 and found 2.18, 1.94, 2.68 and 3.46×10^{25} O/s at $Ls=0^\circ$, 90° , 180° and 270° respectively. This is roughly 6 times smaller than our present estimate. It follows the same dependency with Ls , with a maximum at $Ls=270^\circ$ (perihelion) and a minimum at $Ls=90^\circ$ (aphelion). We discussed this issue in Leblanc et al. (2017) showing that at least a factor 3 of difference can be retrieved by using the same cross sections. We also performed detailed comparison between our two models and with the Fox et al. (2014) calculation of escape probability and did not find significant differences between our two models that might explain the remaining factor 2 difference between them.

The composition of the atmospheric escape induced by sputtering (Figure 4b) is thought to be stoichiometrically related to the composition of the part of the atmosphere where energy deposition of the incident ions occurs (Johnson et al. 2000). Whereas a 1000 eV particle deposits most of its energy around 140 km in altitude, the escape probability for atmospheric particle with few eVs starts to be significant above 150 km in altitude for solar mean conditions. Typically, averaged over the whole atmosphere, the escape probability of a particle with few eV starts to be larger than 0.1 above 160 km at $Ls=0^\circ$ and $Ls=180^\circ$, above 150 km at $Ls=90^\circ$ (aphelion) and 170 km at $Ls=270^\circ$ (perihelion). Considering the escape rates of Ar and N_2 summed with half of the escape rate of the N atoms (because these atoms come from the dissociation of the N_2 molecules impacted by the precipitating particle and by their recoils), we found an excellent agreement (within 10%) between the ratio of these escape rates and the ratio of the average abundances of Ar and N_2 at these altitudes for the four seasons (Chaufray et al. 2014). In other words, the composition of the atmospheric escape

induced by sputtering is close to the composition of the atmosphere at an altitude where the probability of escape of particles with a few eV starts to be larger than 0.1.

IV Conclusion

Using the precipitating flux of heavy ions measured during the first Martian year of MAVEN's science mission, we simulate the expected signatures induced by the sputtering of Mars' upper atmosphere using Exospheric Global Model (Leblanc et al. 2017). Both induced exospheric densities and atmospheric escape are negligible when compared to the equivalent contributions caused by the dissociative recombination of the main ion, O_2^+ , in Mars' ionosphere. The only place where the exospheric density increase caused by sputtering might be observable, is deep in the nightside for the heaviest atmospheric species Ar and CO_2 . However, one of the challenges to identify such signature in MAVEN data is to distinguish an exospheric profile induced by dissociative recombination and sputtering from one induced by thermospheric waves (England et al., 2016).

These results are valid for average conditions and do not preclude the possibility of identifying sputtering signatures when the precipitating flux increases significantly with respect to the average flux shown in this paper (Leblanc et al. 2017). As an example, Hara et al. (2017a and 2017b) have shown that large variation of the precipitating flux can be expected in correlation with the solar wind pressure, the orientation of the electric field of convection and with the presence of crustal magnetic fields.

Extrapolating the present calculation to past solar conditions remains a very uncertain exercise without constraints on the dependency of the precipitating flux intensity with the solar EUV intensity. Such dependency is presently difficult to reconstruct with MAVEN coverage (from solar mean to solar minimum conditions) but should be reachable with MAVEN's second Mars year of operation. However, we can tentatively infer what could be such dependency from Dong et al. (2017) who reported an almost linear dependency of the low energy tailward escaping ion flux with EUV intensity. Using such dependency for the heavy ion precipitation would imply a relatively

modest cumulative total loss of atmosphere via sputtering during the last 4 Gyr (tens of mbar of CO₂ and less than one meter of H₂O equivalent depth of an ocean covering the whole Martian surface) with respect to the expected quantity of atmosphere that should have escaped during that period.

This extrapolation of sputtering induced atmospheric escape loss following a linear dependency with EUV flux is however in contradiction with model (Chaufray et al. 2007) and will need further investigations of MAVEN data to be confirmed or contradicted.

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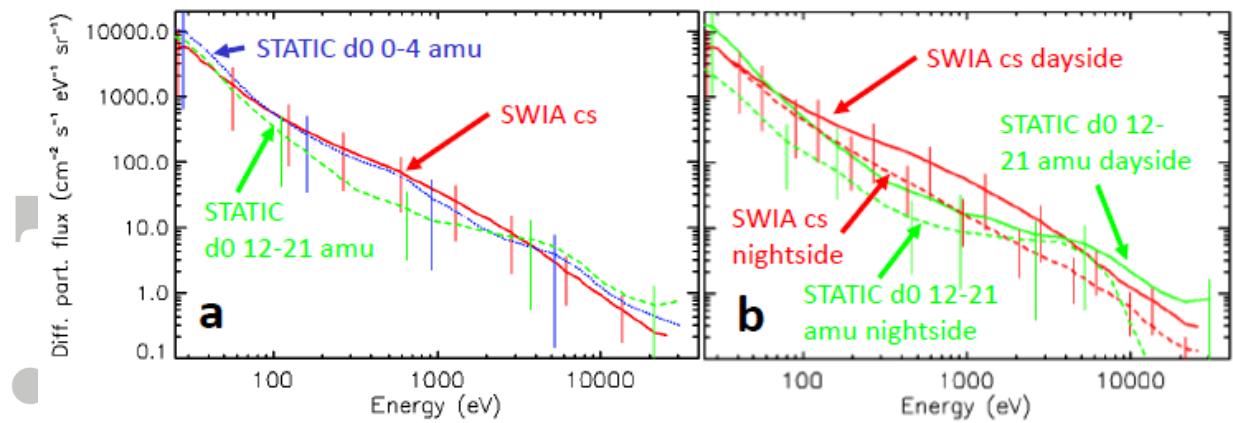


Figure 1: Differential particle flux of the precipitating ion as measured by MAVEN between 200 and 350 km in altitude during the 12/02/2014-10/31/2016 period. Panel a: average flux, solid red line: SWIA cs, dashed green line: STATIC d0 12-21 amu mass range, dotted blue line: STATIC d0 0-4 amu range. Panel b: dayside average flux measured by SWIA (solid red line) and STATIC (d0 product for masses between 12 and 21 amu, in solid green line) and nightside average flux measured by SWIA (dashed red line) and STATIC (d0 product for masses between 12 and 21 amu, in dashed green line). Vertical bars represent the one sigma dispersion of the measurements.

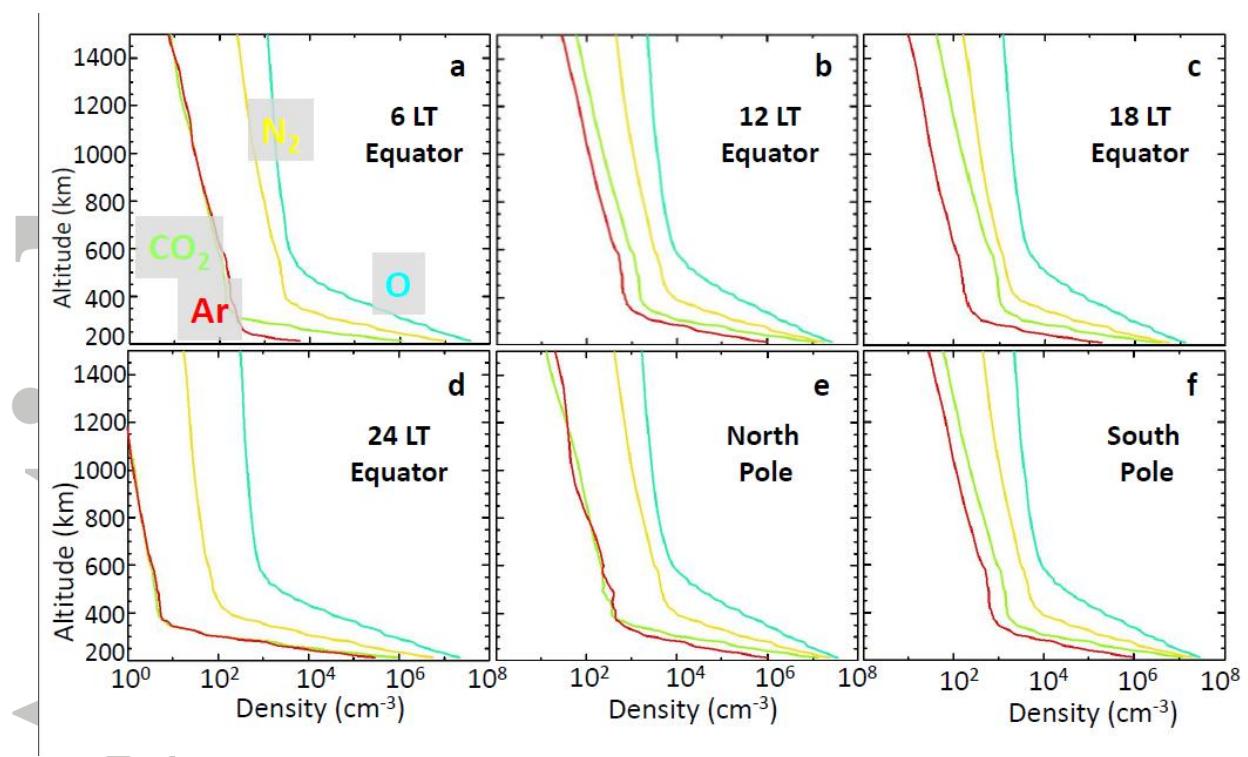


Figure 2: Atmospheric profile (O , CO_2 , N_2 , Ar) at 6 LT equator (panel a), at 12 LT equator (panel b), 18 LT (panel c) equator, 24 LT (panel d) equator, North pole (panel e) and South pole (panel f) from 200 to 1500 km in altitude ($\text{Ls}=180^\circ$, autumn equinox). Blue line: O density, yellow line: N_2 density, green line: CO_2 density, red line: Ar density.

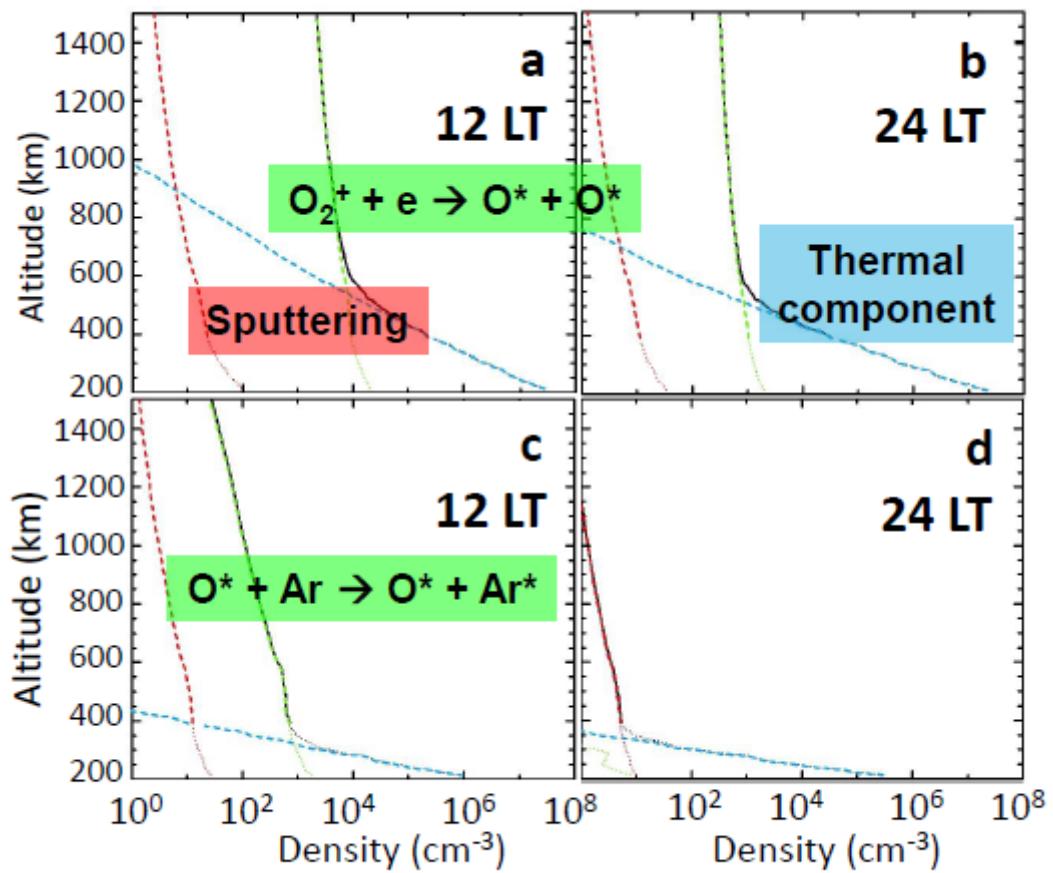


Figure 3: Atmospheric profile of O (panels a and b) and Ar (panels c and d) species at 12 LT (panels a to c) and 24 LT at equator (panels b to d) from 200 to 1500 km in altitude at $Ls=180^\circ$. Sputtering (red dashed line), thermal (blue dashed line) and RD O_2^+ components (green dashed line) and total (dark dashed line).

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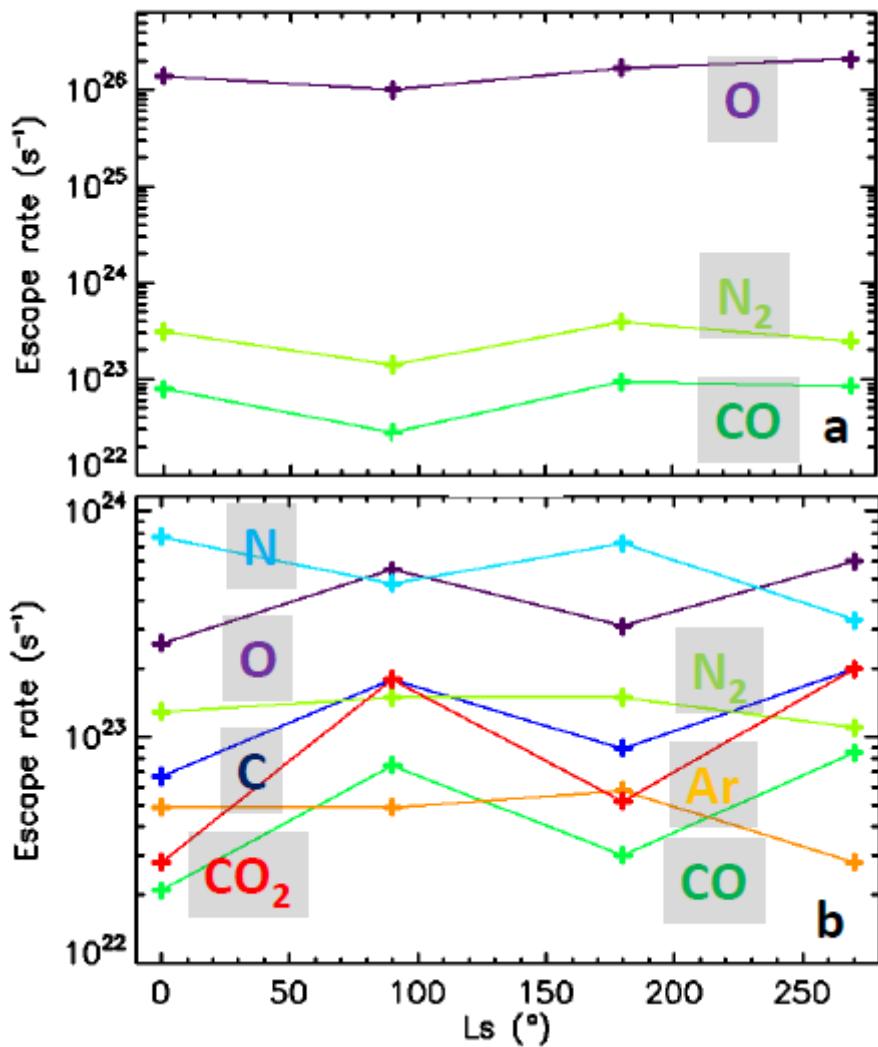


Figure 4: Escape rates of the main heavy martian species at four martian seasons ($Ls=0^\circ$, 90° , 180° and 270°) induced by the dissociate recombination of O_2^+ (panel a) or by sputtering (panel b).