

**What Can We Learn from MHD Simulations and Observations about the Initial Phase of Solar Eruptions?**

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Ideal MHD flux rope instabilities and magnetic (flare) reconnection are believed to be the two main driving mechanisms of solar eruptions, being responsible for the rapid acceleration of the ejecta in such events. However, their onset thresholds and specific trigger mechanisms, as well as their respective contributions to the acceleration, are not well understood from theory, and cannot be inferred unambiguously from observations at present. These limitations greatly hamper our ability to predict the onset and early evolution of individual eruptions. In this talk, I will briefly discuss to what extent MHD simulations and their comparison to observations can be utilized to improve our understanding of the initial phase of solar eruptions.